

The little-studied cluster Berkeley 90

II. The foreground ISM

J. Maíz Apellániz¹, R. H. Barbá², A. Sota³, and S. Simón-Díaz^{4,5}

¹ Centro de Astrobiología, CSIC-INTA, campus ESAC, apartado postal 78, E-28 691 Villanueva de la Cañada, Madrid, Spain
e-mail: jmaiz@cab.inta-csic.es

² Departamento de Física y Astronomía, Universidad de La Serena, Av. Cisternas 1200 Norte, La Serena, Chile

³ Instituto de Astrofísica de Andalucía-CSIC, Glorieta de la Astronomía s/n, E-18 008 Granada, Spain

⁴ Instituto de Astrofísica de Canarias, E-38 200 La Laguna, Tenerife, Spain

⁵ Departamento de Astrofísica, Universidad de La Laguna, E-38 205 La Laguna, Tenerife, Spain

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ABSTRACT

Context. Nearly one century after their discovery, the carrier(s) of Diffuse Interstellar Bands is/are still unknown and there are few sightlines studied in detail for a large number of DIBs.

Aims. We want to study the ISM sightlines towards LS III +46 11 and LS III +46 12, two early-O-type stellar systems, and LS III +46 11 B, a mid-B-type star. The three targets are located in the stellar cluster Berkeley 90 and have a high extinction.

Methods. We use the multi-epoch high-S/N optical spectra presented in Paper I (Maíz Apellániz et al. 2015), the extinction results derived there, and additional spectra.

Results. We have measured equivalent widths, velocities, and FWHMs for a large number of absorption lines in the rich ISM spectrum in front of Berkeley 90. The absorbing ISM has at least two clouds at different velocities, one with a lower column density (thinner) in the K I lines located away from Berkeley 90 and another one with a higher column density (thicker) associated with the cluster. The first cloud has similar properties for both O-star sightlines but the second one is thicker for LS III +46 11. The comparison between species indicate that the cloud with a higher column density has a denser core, allowing us to classify the DIBs in a $\sigma - \zeta$ scale, some of them for the first time. The LS III +46 12 sightline also has a high-velocity redshifted component.

Key words. Dust, extinction — ISM: lines and bands — Open clusters and associations: individual: Berkeley 90 — Stars: early-type — Stars: individual: LS III +46 11 — Stars: individual: LS III +46 12

1. Introduction

Diffuse interstellar bands (DIBs) were discovered nearly a century ago (Heger 1922; Merrill 1934). There are many of them known, with FWHMs that range from under 1 Å to several tens of Å. They are correlated with the amount of extinction but not perfectly so. To date, the carrier or carriers that produce them have not been identified with certainty. Different options have been proposed, most of them carbon-based substances (Herbig 1995; Ehrenfreund et al. 1995; Thorburn et al. 2003; Kaźmierczak et al. 2010; Maier et al. 2011; Salama et al. 2011; Steglich et al. 2011).

Broadly speaking, DIB observational studies can be classified in two categories. Those in the first one select one or a small number of DIBs and measure them for a large sample of stars. The pioneering study was Duke (1951) and in recent years this type has become popular (Thorburn et al. 2003; Munari et al. 2008; Friedman et al. 2011; Vos et al. 2011; Raimond et al. 2012; Puspitarini et al. 2013; van Loon et al. 2013). The second category corresponds to studies that select one or a small number of stars and study their DIBs in depth throughout a large wavelength range (Jenniskens & Desert 1994;

Galazutdinov et al. 2000; Tuairisg et al. 2000; Cox et al. 2005; Hobbs et al. 2008, 2009). The work presented here is of the second category, though it is part of a larger project that plans to study the ISM using thousands of high-quality spectra of OB stars (Penadés Ordaz et al. 2013; Maíz Apellániz et al. 2014b).

The sparse young open cluster Berkeley 90 is almost unstudied. In Maíz Apellániz et al. (2015b), from now on Paper I, we analyzed it with an emphasis on LS III +46 11, a very massive early-type binary composed of two of the four O stars earlier than O4 known in the northern hemisphere. Paper I also studied LS III +46 12, the other early-O-type system with a significant contribution to the ionizing flux of the cluster, and other properties of Berkeley 90. The cluster is immersed in an H II region, Sh 2-115. In Table 1 we summarize some of the properties of the two dominant systems in the cluster and of a third star that is also studied here. Note that the two dominant stars experience a similar extinction law, as evidenced by the measured R_{5495} values (Maíz Apellániz 2013a; Maíz Apellániz et al. 2014a), but LS III +46 11 has an extinction which is higher by $\approx 30\%$. In this paper we study the signature of the foreground ISM in the optical spectra of both stars. We also include in our study a third system, LS III +46 11 B, which is spatially resolved from

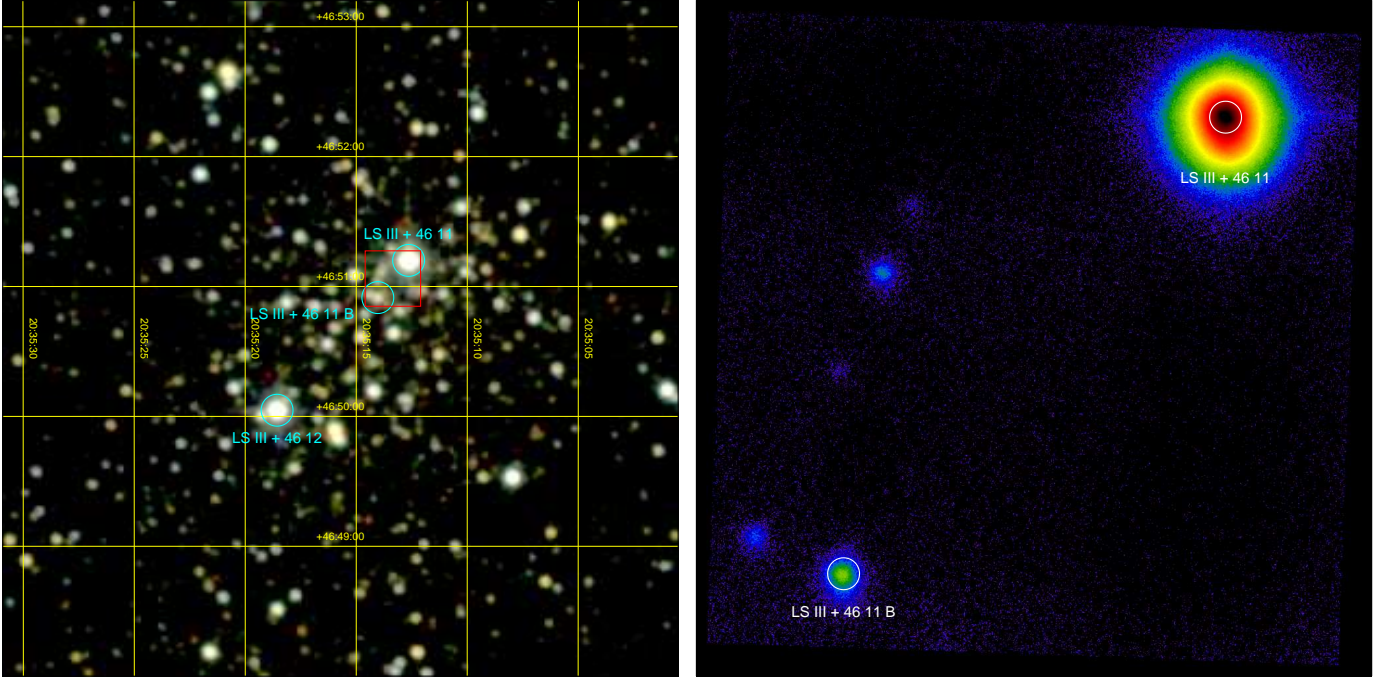


Fig. 1. [left] 2MASS K_sHJ three-color RGB mosaic of Berkeley 90. The intensity level in each channel is logarithmic. [right] False-color AstraLux $25'6 \times 25'6$ i -band image of LS III +46 11 (see Paper I). North is up and East is left and the image location corresponds to the red square in the left panel. The intensity level is logarithmic.

	LS III +46 11	LS III +46 12	LS III +46 11 B	Reference
Sp. type	O3.5 If* + O3.5 If*	O4.5 V((f))	B4 V	Paper I, this work
RA (J2000)	20:35:12.642	20:35:18.566	20:35:14.029	Høg et al. (2000)
dec (J2000)	+46:51:12.12	+46:50:02.90	+46:50:54.94	Høg et al. (2000)
l (deg.)	84.8844	84.8791	84.8830	Høg et al. (2000)
b (deg.)	+3.8086	+3.7836	+3.8026	Høg et al. (2000)
V_J	10.889 ± 0.021	10.268 ± 0.009	$15.992 \pm 0.033^*$	Paper I, this work
K_s	6.971 ± 0.023	7.470 ± 0.023	11.821 ± 0.053	Skrutskie et al. (2006), Paper I
R_{5495}	3.303 ± 0.058	3.377 ± 0.040	3.34 (fixed)	Paper I, this work
$E(4405 - 5495)$	1.653 ± 0.020	1.255 ± 0.011	1.590 ± 0.019	Paper I, this work
A_V	5.475 ± 0.037	4.272 ± 0.021	5.324 ± 0.064	Paper I, this work
v (km/s)	-20 ± 5	-13 ± 3	—	Paper I

* Calculated from CHORIZOS, not measured.

Table 1. LS III +46 11, LS III +46 12, and LS III +46 11 B summary.

LS III +46 11 (Fig. 1) and whose photometry was presented in Paper I, for which we have also obtained spectra.

2. Data

Most of the spectra used here are a selection of those used in Paper I. The reader is referred to that article for details; here we only summarize the most relevant information. The spectra were obtained under four different projects (GOSSS, Maíz Apellániz et al. 2011; NoMaDS, Maíz Apellániz et al. 2012, Pellerin et al. 2012; IACOB, Simón-Díaz et al. 2011, Simón-Díaz et al. 2015; and CAFÉ-BEANS, Negueruela et al. 2015) at resolving powers ranging from 2800 to 65 000 and covering different ranges from 3811 Å to 9225 Å.

In addition to the spectra from Paper I, we have obtained long-slit spectroscopy on 29 April 2015 with the OSIRIS instrument at the 10.4 m Gran Telescopio Canarias (GTC), as part of an extension of GOSSS to larger-aperture telescopes to include fainter O stars. Two gratings were used, R2500U (3440–4610 Å, t_{exp} of 1800 s) and R2500V (4445–6060 Å, t_{exp} of 300 s), with

a resolving power R between 2500 and 3000. As in the rest of the GOSSS data, the long slit was placed to include both LS III +46 11 and LS III +46 12. The larger aperture of GTC allowed us to extract the spectra of a third star, LS III +46 11 B (2MASS J20351402+4650549, see Paper I and Fig. 1), that was present in the long-slit exposures.

Note that the number of epochs, S/N, and wavelength coverage at each resolution for each star varies, so there are differences in the ISM lines that can be studied for each star. In particular, the blue-violet range was significantly less covered for LS III +46 12 than for LS III +46 11 and the number of lines that could be measured for LS III +46 11 B was relatively small. The telluric lines in the high-resolution spectra were eliminated according to the procedure of Gardini et al. (2013).

The results in this paper can be divided into two types. We first conduct a brief analysis of the spectra and photometry of LS III +46 11 B in order to derive its properties and place it in the context of Berkeley 90. We then study the ISM in front of the three targets, which is the core of this paper.

Type	λ_0 (Å)	EW (mÅ)		FWHM (Å)		v km/s		Kf	Mf		Ref.	Notes
		S1	S2	S1	S2	S1	S2		S1	S2		
CN	3873.999	25.2±1.8	—	0.123±0.010	—	-18.2±0.4	—	-	HF	-	4	-
CN	3874.607	48.1±1.9	—	0.144±0.007	—	-17.8±0.2	—	-	HF	-	4	-
CN	3875.760	14.4±2.8	—	0.132±0.020	—	-17.2±0.7	—	-	HF	-	4	-
CH	3886.410	22.2±2.8	—	0.135±0.019	—	-20.6±0.9	—	-	HF	-	4	-
Ca II	3933.663	390.8±7.1	509.7±4.0	—	—	-15.0±0.5	35.0±1.4	M	HI	GI	5	-
CH+	3957.692	29.4±1.8	—	0.188±0.013	—	-18.8±0.6	—	-	HF	-	4	-
Ca II	3968.468	290.4±3.7	—	—	—	-15.0±0.2	—	M	HI	-	5	-
Ca I	4226.7275	13.9±1.1	—	0.188±0.016	—	-20.1±0.8	—	-	HF	-	5	-
CH+	4232.548	46.0±1.5	50.0±1.8	—	—	-17.8±0.2	-19.5±1.7	M	HI	GI	4	-
CH	4300.313	75.3±1.3	41.2±1.3	—	—	-17.3±0.5	-23.1±1.3	M	HI	GI	4	-
DIB	4427.94	1641.4±23.5	1795.6±19.4	17.333±0.104	17.270±0.115	15.3±3.1	14.9±3.1	-	GF	GF	3	-
DIB	4501.67	216.3±6.9	177.7±8.5	—	—	13.2±1.9	-16.5±1.7	-	GI	GI	3	1
DIB	4726.70	207.2±2.3	179.5±2.5	—	—	7.3±2.1	-4.1±1.6	-	GI	GI	3	1
DIB	4761.12	413.0±13.1	364.4±21.3	22.000±1.000	22.000±1.000	195.9±21.7	126.7±23.9	-	GF	GF	3	6
DIB	4762.36	87.4±4.1	99.1±4.1	2.69	2.69	-2.0±4.2	-17.2±3.7	-	GF	GF	3	5
DIB	4779.69	69.3±3.9	66.3±3.9	3.000±0.300	3.000±0.300	22.3±6.0	-1.0±6.3	-	GF	GF	3	4
DIB	4879.83	196.0±12.5	206.5±17.7	11.51	11.51	-13.6±13.6	-116.7±24.9	-	GF	GF	3	8
DIB	4887.43	1128.5±136.4	1453.4±104.1	39.75	39.75	-13.6±13.6	-116.7±24.9	-	GF	GF	3	7
DIB	4963.85	42.4±1.5	—	0.659±0.027	—	-14.0±0.4	—	-	HF	-	3	-
DIB	4984.59	17.4±1.2	—	0.549±0.032	—	-6.7±0.8	—	-	HF	-	3	-
DIB	5236.29	27.4±2.6	—	1.561±0.099	—	-15.5±1.6	—	-	HF	-	3	-
DIB	5245.43	102.1±5.2	—	7.26	—	-2.6±9.2	—	-	HF	-	3	-
DIB	5363.52	35.5±3.9	—	2.228±0.224	—	-7.4±2.5	—	-	HF	-	3	-
DIB	5404.56	17.4±1.1	23.7±3.1	0.816±0.036	0.999±0.120	-20.5±0.5	-15.8±2.3	-	HF	HF	1	-
DIB	5418.87	23.7±0.9	20.9±3.9	0.741±0.019	0.811±0.120	-17.3±0.4	-14.5±2.2	-	HF	HF	1	-
DIB	5449.83	301.1±11.7	288.8±11.7	14.06	14.06	-46.5±7.7	-54.9±6.8	-	GF	GF	3	-
DIB	5487.23	147.1±6.9	175.5±8.9	4.776±0.156	5.017±0.171	-1.7±2.8	-17.2±2.4	-	GF	GF	3	-
DIB	5494.29	32.7±0.8	34.9±2.3	0.705±0.022	0.699±0.048	-28.7±0.3	-25.7±1.0	-	HF	HF	3	-
DIB	5506.28	22.4±0.9	—	1.30	—	-20.9±1.4	—	-	HF	-	1	-
DIB	5508.12	61.0±3.1	—	1.899±0.042	—	-21.3±0.9	—	-	HF	-	1	-
DIB	5512.68	14.6±0.5	—	0.71	—	-15.9±2.1	—	-	HF	-	1	-
DIB	5545.06	31.1±2.6	—	0.921±0.051	—	-16.6±0.9	—	-	HF	-	1	10
DIB	5705.08	146.3±1.3	168.6±5.2	3.70	3.70	-12.2±0.7	-6.5±2.8	-	HF	HF	1	-
DIB	5780.48	496.8±1.4	563.6±3.9	—	—	-11.3±0.2	-10.3±0.5	-	HI	HI	1	1
DIB	5797.06	167.3±1.4	155.9±3.2	—	—	-12.7±0.3	-12.7±0.8	-	HI	HI	1	1
DIB	5849.81	61.9±0.6	59.4±1.6	0.82	0.82	-17.2±0.2	-18.0±0.6	-	HF	HF	1	-
Na I	5889.951	593.3±3.3	667.6±8.5	—	—	-14.2±0.1	2.8±1.2	M	HI	HI	5	-
Na I	5895.924	553.1±3.4	607.1±7.6	—	—	-14.3±0.1	-2.1±1.2	M	HI	HI	5	-
DIB	6010.75	115.4±1.4	130.6±4.0	3.27	3.27	-31.9±1.0	-34.6±2.4	-	HF	HF	1	11
DIB	6113.18	20.0±0.4	17.7±1.1	0.68	0.68	-18.2±0.4	-17.6±1.1	-	HF	HF	1	-
DIB	6139.98	10.5±0.6	—	0.469±0.041	—	-20.1±0.9	—	-	HF	HF	1	-
DIB	6177.30	1338.6±54.8	1546.0±17.0	23.06	23.06	-105.9±12.3	-103.6±5.4	-	HF	HF	2	2
DIB	6195.98	56.3±0.3	58.5±1.1	0.42	0.42	-17.7±0.1	-17.6±0.4	-	HF	HF	1	12
DIB	6203.05	94.3±1.0	96.7±3.0	1.20	1.20	-18.5±0.2	-17.7±0.7	-	HF	HF	1	3
DIB	6269.85	86.1±0.7	96.4±2.1	1.18	1.18	-20.4±0.2	-21.2±0.6	-	HF	HF	1	9
DIB	6283.84	1037.0±17.3	1182.0±22.8	—	—	8.7±1.8	12.8±6.2	-	HI	HI	1	1
DIB	6379.32	70.0±1.3	62.3±3.8	0.59	0.59	-20.1±0.1	-21.2±0.7	-	HF	HF	1	-
DIB	6425.66	17.1±0.4	16.9±1.3	0.75	0.75	-16.0±0.4	-14.3±1.5	-	HF	HF	1	-
DIB	6439.48	18.0±0.3	17.2±1.2	0.75	0.75	-17.4±0.6	-13.0±1.4	-	HF	HF	1	-
DIB	6445.28	26.5±0.7	30.9±1.9	—	—	-17.4±0.3	-16.4±1.3	-	HI	HI	1	1
DIB	6449.22	21.0±0.8	—	0.823±0.045	—	-18.5±0.7	—	-	HF	HF	1	-
DIB	6456.01	31.0±0.6	35.4±3.0	—	—	-18.2±0.5	-20.6±1.6	-	HI	HI	1	1
DIB	6520.62	19.0±0.5	—	—	—	-15.4±0.5	—	-	HI	-	1	1
DIB	6590.42	412.5±19.1	343.4±23.0	12.823±0.248	10.402±0.714	14.8±3.4	50.5±7.8	-	HF	HF	1	-
DIB	6597.31	12.4±0.4	13.6±2.1	0.54	0.54	-16.1±0.5	-16.8±2.1	-	HF	HF	1	-
DIB	6613.62	237.3±0.9	246.0±5.8	—	—	-12.7±0.1	-11.4±0.5	-	HI	HI	1	1
DIB	6660.71	29.2±0.3	32.1±1.1	0.59	0.59	-19.9±0.2	-18.3±0.5	-	HF	HF	1	-
DIB	6672.27	17.1±0.5	18.7±1.5	—	—	-17.9±0.5	-22.0±1.1	-	HI	HI	1	1
DIB	6699.32	24.1±0.5	21.5±1.5	0.64	0.64	-19.7±0.3	-19.4±1.1	-	HF	HF	1	13
DIB	6993.13	86.3±0.4	90.8±3.0	0.75	0.75	-18.1±0.1	-18.5±0.5	-	HF	HF	1	-
DIB	7224.03	225.9±1.3	253.2±3.3	—	—	-14.0±0.2	-13.3±0.5	-	HI	HI	1	1
K I	7664.911	390.1±4.9	390.5±14.2	—	—	-15.1±0.1	-14.8±0.4	M	HI	HI	5	-
K I	7698.974	299.8±1.5	250.3±3.8	—	—	-15.7±0.1	-16.1±0.2	M	HI	HI	5	-
DIB	8621.20	305.3±12.9	364.5±20.0	4.688±0.127	3.684±0.236	-36.3±2.0	-36.6±2.4	-	HF	HF	2	-

References: 1: Hobbs et al. (2008). 2: Jenniskens & Desert (1994). 3: Maíz Apellániz et al. (2014b). 4: Smoker et al. (2014). 5: van Hoof (1999).

Notes: 1: Asymmetric profile. 2: Broad and asymmetric but double fit not attempted. 3: Inside another broader unmeasured DIB, parabolic background used. 4: Multiple DIB with 4761.12 Å and 4762.36 Å. 5: Multiple DIB with 4761.12 Å and 4779.69 Å. 6: Multiple DIB with 4762.36 Å and 4779.69 Å. 7: Multiple DIB with 4879.83 Å. 8: Multiple DIB with 4887.43 Å. 9: Possibly inside broader DIB. 10: Weaker DIB at 5546.08 Å not fit. 11: Weaker DIB at 6004.89 Å not fit. 12: Weaker DIB at 6194.74 Å not fit. 13: Weaker DIB at 6702.02 Å not fit.

Table 2. ISM lines measured for LS III +46 11 (S1) and LS III +46 12 (S2) in the WHT and high-resolution spectra. Empty values for the EW and v indicate that no good-quality spectra were available for LS III +46 12. A single value for the FWHM indicates that a reference value was used and left fixed, one with an uncertainty indicates that it was fitted in each spectrum, and an empty value that Gaussian fitting was not used or no good-quality spectra were available. Each line was fitted individually except for the two DIB groups at 4761.12+4762.36+4779.69 Å and 4879.83+4887.43 Å (Maíz Apellániz et al. 2014b), which were fitted simultaneously. Kf stands for kinematic multiplicity flag (one Gaussian component can be measured at each of the two cloud velocities, see text). Mf is a measurement type flag with two components: [a] H indicates that the high-resolution data were used while G indicates that the GOSSS WHT data were used and [b] F indicates that a Gaussian fit was used while I indicates that numerical integration was used instead. Ref. gives the reference used for the rest wavelength (λ_0) and for the FWHM (when fixed).

For the ISM we have measured as many as seven atomic and seven molecular lines and fifty DIBs. We give the equivalent widths (EWs), full-widths at half maximum (FWHMs), and central velocities (v) for LS III +46 11 and LS III +46 12 obtained with the WHT and the high-resolution spectroscopy in Table 2. Before analyzing the results, we detail some aspects of the extraction:

- Our spectra are diverse in terms of wavelength coverage, S/N, and spectral resolution. Therefore, each absorption line was extracted spectrum by spectrum and the results were combined a posteriori using the best available data in each case.
- In cases where spectra of similar resolutions but different S/N values were available, we only used those where the S/N was adequate.
- For narrow lines, high-resolution spectroscopy is preferred over intermediate-resolution data. For broad DIBs, however, the situation is reversed as measurement errors are dominated by normalization, which is easier to carry out in long-slit spectra than in echelle data. For the intermediate-resolution case we only used the WHT spectra due to their slightly higher spectral resolution compared to the CAHA-3.5 m data.
- In most cases, fitting with a single Gaussian yields better results than numerical integration and the former are the ones displayed in Fig. 2. However, some lines are asymmetric due to their intrinsic profiles (Galazutdinov et al. 2008) or the existence of two or three kinematic components (see below). For those lines we display the numerical integration results in Fig. 2.
- No attempt was done to fit broad DIBs with Lorentzian profiles instead of Gaussians (Snow et al. 2002; van Loon et al. 2013) because of the rectification difficulties that exist for the wings. Note, however, that differences in the results are small (Maíz Apellániz et al. 2014b).
- For the DIB groups at 4761.12+4762.36+4779.69 Å and 4879.83+4887.43 Å (Maíz Apellániz et al. 2014b), the lines were fitted simultaneously. Note, however, that for the DIB at 6203.05 Å only the inner narrow DIB was fit (but not the broad one) and that a single Gaussian was also used for the 6177.30 Å DIB despite its obvious asymmetry due to the complexity of the spectral region¹.
- The reference values and sources for the rest wavelength (λ_0) of each line are given in Table 2. For the DIBs with Gaussian fits we first attempted fixing the FWHM using the reference value from the source. If that first attempt was successful, it is listed in the table as a single value. If we thought that the FWHM could be improved, we did a second attempt leaving it as a free parameter for each fit. Those cases are listed with a measurement and an uncertainty in Table 2 and those with the most significant changes also appear in Table 3.

The EWs obtained from the GTC spectra for the three targets are given in Table 4. We followed the same procedure as for the WHT and high-resolution spectroscopy with only a few differences:

- We only measured those lines where additional information could be added to the one in Table 2.

¹ The velocities for the 6177.30 Å DIB differ from the rest of the DIBs precisely for this reason and should not be considered; only its EWs in Table 2 are of practical use. Nevertheless, note that velocities for LS III +46 11 and LS III +46 12 are consistent between them, a sign that they originate in the same clouds.

λ_0 (Å)		FWHM (Å)	
old	new	old	new
4427.94	4428.42	24.15	17.30
4761.12	4765.0	19.72	22.00
	4779.69	5.48	3.00
	4963.85	2.62	0.66
	4984.59	1.33	0.55
	5236.29	2.27	1.56
5449.83	5451.1		14.06
	5487.23	6.63	4.78
	5494.29	1.90	0.71
6010.75	6010.43		3.27
6283.84	6284.36		4.77
6590.42	6591.5	7.53	12.82
8621.20	8620.65	1.86	4.69

Table 3. DIBs for which the measured values of λ_0 and/or FWHM are significantly different for LS III +46 11 and LS III +46 12 when compared with the reference values. A single value indicates no apparent change. The number of significant digits in λ_0 reflects the precision of the value, which is worse for weak/broad DIBs than for strong/narrow ones. The values listed here originate from both the high- and intermediate-resolution spectroscopy (see Table 2) but the latter is not used for the narrow DIBs to avoid resolution issues.

- For LS III +46 11 B we extracted the absorption lines after subtracting a model TLUSTY SED (see below). This allowed us to measure Ca II λ 3968.468 for that sightline (but not for the other two).
- The three CN lines at 3873.999+74.607+75.760 could not be individually resolved so we give the total EW.
- For CH+ λ 3957.692 and Ca I λ 4226.7275 we would not obtain an EW for LS III +46 11 B due to the poor S/N.
- A single measurement flag (Mf) is used to indicate either a Gaussian fit (F) or numerical integration (I) was used to obtain the EW.
- In most cases the uncertainties for LS III +46 11 and LS III +46 12 are larger than in Table 2 despite the larger aperture of GTC compared to the other telescopes. This is caused by the lower spectral resolution, fewer epochs, and shorter exposure times.

3. LS III +46 11 B

The blue-violet section of the GTC combined spectrum of LS III +46 11 B is shown at the top of Fig. 2. The jump in S/N at 4590 Å is caused by the use of two gratings, R2500U and R2500V, to the left and right, respectively, with different exposure times.

As we did in Paper I for LS III +46 11 and LS III +46 12, we used MGB (Maíz Apellániz et al. 2012) to derive the stellar properties of LS III +46 11 B. We built a classification grid based not on observed spectra but on the TLUSTY grids of Lanz & Hubeny (2003, 2007) using their SYNSPEC output. The grid uses T_{eff} (15 000–55 000 K) as a horizontal coordinate and $\log g$ (1.75–4.75 cgs) as vertical coordinate and, as before, MGB allows for the variation of the rotation speed $v \sin i$.

MGB yields a good fit to the observed spectrum with $T_{\text{eff}} = 16\,000 \pm 1000$ K, $\log g = 4.00 \pm 0.25$ cgs, and $v \sin i = 175 \pm 25$ km/s. LS III +46 11 B appears to be a normal main-sequence

Type	λ_0 (Å)	EW (mÅ)			Mf
		S1	S1B	S2	
CN	3873.999+74.607+75.760	87.4±6.4	115.1±49.0	8.7±1.2	F
Ca II	3933.663	382.6±6.5	396.1±41.3	528.1±8.4	I
CH+	3957.692	43.8±9.4	—	31.7±5.0	F
Ca II	3968.468	—	252.8±31.8	—	I
Ca I	4226.7275	15.2±3.6	—	15.9±3.1	F
CH+	4232.548	45.2±2.5	43.0±10.8	49.1±3.6	F
CH	4300.313	74.9±2.1	71.1±4.4	43.8±2.0	F
DIB	4427.94	1652.3±13.6	1636.7±75.7	1784.7±12.2	F
DIB	4501.67	208.1±4.9	186.3±21.6	170.7±11.5	I
DIB	5487.23	160.3±23.9	151.4±40.5	177.4±16.2	F
DIB	5705.08	158.6±13.2	166.3±32.1	160.6±10.0	F
DIB	5780.48	508.4±8.8	526.4±22.3	570.9±7.4	I
DIB	5797.06	172.0±8.8	201.6±30.7	176.8±31.6	I
Na I	5889.951	593.8±7.8	550.1±47.7	655.9±10.4	I
Na I	5895.924	555.9±20.2	559.4±44.8	624.1±23.1	I

Table 4. EWs measured for LS III +46 11 (S1), LS III +46 11 B (S1B), and LS III +46 12 (S2) in the GTC spectra. Empty values indicate that no measurement could be obtained. In the Mf column, F indicates that a Gaussian fit was used while I indicates that numerical integration was used instead.

Quantity	Baseline	Alt. LC	Alt. T_{eff}
T_{eff} (K)	16 000	16 000	17 000
luminosity class	5.5	5.0	5.5
χ^2_{red}	0.50	0.48	0.51
R_{5495}	3.34	3.34	3.34
$E(4405 - 5495)$ (mag)	1.590±0.019	1.590±0.019	1.606±0.019
A_V (mag)	5.324±0.064	5.323±0.064	5.376±0.063
$V_{J,0}$ (mag)	10.668±0.040	10.669±0.040	10.620±0.040
$\log d$ (pc)	3.197±0.008	3.335±0.008	3.223±0.008

Table 5. Results of the CHORIZOS fits for LS III +46 11 B.

mid-B star² with a moderately high rotation speed. The value of $\log g$ is slightly lower than the expected one for a $\sim 4.5 M_{\odot}$ star of the young age of Berkeley 90 but the discrepancy is of only 1 sigma. Note that the estimated uncertainties are obtained by eye comparing different models in the grid i.e. they are not formal uncertainties derived by e.g. χ^2 fitting. The best TLUSTY fit is shown in the middle of Fig. 2 and the residual of the fit at the bottom. The residual yields the ISM spectrum used to measure the EWs of its absorption lines. In particular, note how the residual has subtracted H ϵ correctly, leaving the narrow ISM Ca II $\lambda 3968.468$ line clearly visible.

We have also done a CHORIZOS (Maíz Apellániz 2004) analysis of LS III +46 11 B, similar to those in Paper I for the O stars but with some differences:

- As input photometry we used 2MASS (JHK_S , Skrutskie et al. 2006) and IPHAS (ri , Barentsen et al. 2014), as there is no other high-quality optical data available. As it happened for LS III +46 11 with K_S , there is no good detection of LS III +46 11 B in the 2MASS J photometry. We downloaded the image from the archive and we performed a differential photometry analysis similar to the one we did for LS III +46 11 in Paper I. That yielded a J magnitude of 12.552 ± 0.050 for LS III +46 11 B.

- We used the same Milky Way grid (Maíz Apellániz 2013b) and extinction laws (Maíz Apellániz 2013a; Maíz Apellániz et al. 2014a) as in Paper I.
- For the baseline CHORIZOS run we fixed T_{eff} to 16 000 K (see above), the photometric luminosity class (LC) to 5.5 (ZAMS, since the star is expected to be ~ 2 Ma, young for the lifetime of a mid-B star in the main sequence), and the extinction type (R_{5495}) to 3.34 (an average of the two very similar values measured in Paper I for LS III +46 11 and LS III +46 12). The amount of extinction [$E(4405 - 5495)$] and logarithmic distance ($\log d$) were left as free parameters.
- We did two additional runs, one changing LC to 5.0 (average main sequence) and another changing T_{eff} to 17 000 K to quantify the effect of the parameter uncertainty in the output.

The CHORIZOS results are shown in Table 5:

- The three runs give similar results for χ^2_{red} and $E(4405 - 5495)$ and only differ significantly in $\log d$. This happens because the optical+NIR colors to the right of the Balmer jump for early/mid B stars are nearly independent of luminosity and only weakly dependent on T_{eff} .
- The values of χ^2_{red} indicate that the fit is very good in all cases (due to the reasons in the previous point they are all expected to be similar). This is in principle a sign of the validity of the extinction laws and the R_{5495} value but it is only a weak sign, as the SED is not probed for $\lambda < 6000$ Å.
- The LS III +46 11 B amount of extinction is intermediate between those of LS III +46 11 and LS III +46 12, but significantly closer to that of LS III +46 11 (which is a shorter distance away in the plane of sky).
- If LS III +46 11 B is close to the ZAMS, $\log d$ must be close to 3.20. Allowing for an older age places the system slightly farther away but not as far as the preferred value of 3.40–3.45 derived in Paper I, which required that LS III +46 12 is a hidden binary. A $\log d$ close to 3.20 poses no problem for a single LS III +46 12 but it implies that, if all the stars are at the same distance, the two spectroscopic components of LS III +46 11 would have similar visual luminosities as LS III +46 12, which contradicts their supergiant spectroscopic classification. An alternative, already hinted in Paper

² Comparison with our original GOSSS standard grid extended to B stars using MGB yields a B4 V spectral type (Maíz Apellániz et al. 2015a).

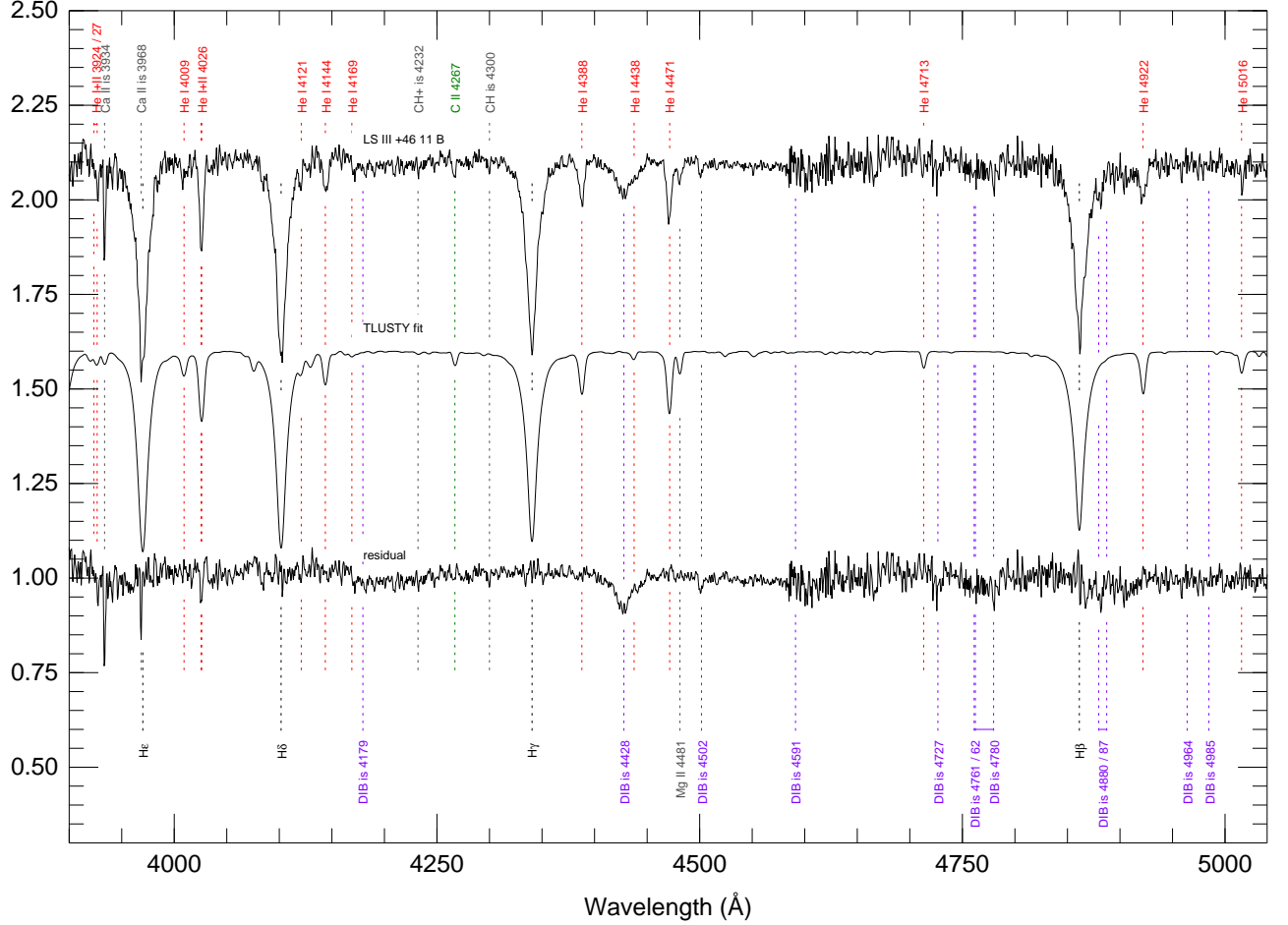


Fig. 2. [top] GTC combined spectrum of LS III +46 11 B. [middle] Best TLUSTY fit for LS III +46 11 B. [bottom] Residual obtained by dividing the top spectrum by the middle fit.

I, would imply that He II $\lambda 4685.71$ for the earliest O stars is a measurement of wind strength but not of luminosity.

4. ISM results

We analyze the ISM lines in three steps. We first concentrate on the velocities in Table 2 and the EWs in Tables 2 and 4:

- All the EWs with values in both Tables 2 and 4 are consistent within the errors, so from now on we will refer only to the most precise value of the two in each case.
- All the atomic and molecular lines for LS III +46 11 have velocities between -20.6 and -13.8 km/s. CH+ $\lambda 4232.548$, CH $\lambda 4300.313$, and K I $\lambda 7664.911 + 7698.974$ for LS III +46 12 also have similar velocities. However, the Ca II $\lambda 3933.663$ and Na I $\lambda 5889.951 + 5895.924$ velocities for that star are significantly larger, especially in the first case. The cause is explained below.
- The two R(0) lines of CH+ are easily detected in the NoMaDS spectra of LS III +46 11 but the R(1) and Q(1) lines are not, possibly due to an insufficient S/N (Oka et al. 2013).
- The majority of the DIBs for both stars have velocities close to those of the atomic and molecular lines, indicating that the DIBs originate at the same clouds (from one point of view) or that we have a good knowledge of their central wavelengths (from another point of view). The most

clear exceptions are listed in Table 3, where we computed the new values of λ_0 assuming that the DIBs originated at the same velocity as the atomic and molecular lines. The majority of the changes listed there correspond to values measured by Maíz Apellániz et al. (2014b), who used only intermediate-resolution spectroscopy (which biases FWHM towards larger values) and combined profiles from different stars (which can also broaden the FWHM as well as introduce biases in λ_0). Among those cases we find the well known DIB with $\lambda \approx 4428$ Å, which is very strong but broad and contaminated by weak stellar lines: the value found here is ≈ 0.5 Å towards the blue compared with the one in Maíz Apellániz et al. (2014b). The 8621 Å DIB was not included on the Hobbs et al. (2008) or Maíz Apellániz et al. (2014b) lists so the reference values listed in Tables 2 and 3 are from Jenniskens & Desert (1994). This DIB has received recent attention due to its inclusion in large-scale surveys such as RAVE and Gaia, so it is important to know its intrinsic characteristics with accuracy. Our values for λ_0 and FWHM are significantly different from the reference ones but we should note that more recent studies (e.g. Munari et al. 2008; Hobbs et al. 2009; Kos et al. 2013) have also pointed towards changes in the same direction (lower central wavelengths and broader profiles).

- The EWs for the LS III +46 11 B lines have, in general, relatively large errors. However, they are all consistent with hav-

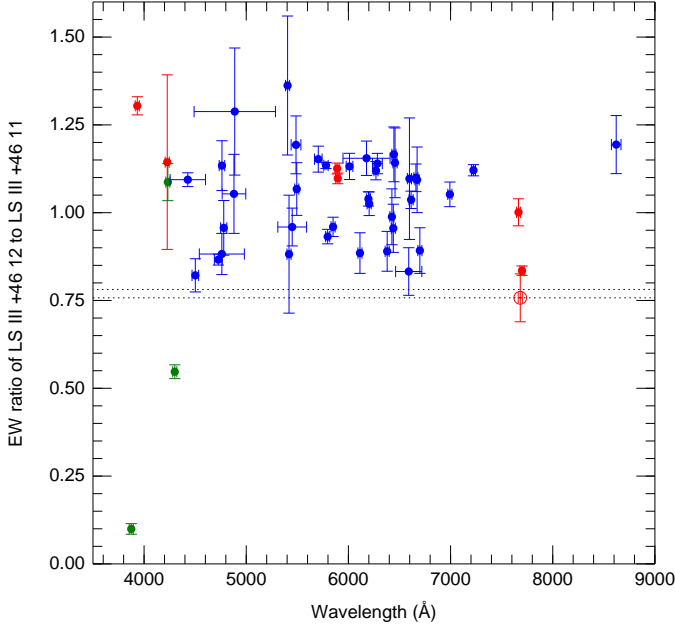


Fig. 3. Equivalent-width ratios between LS III +46 12 and LS III +46 11 for the cases in Table 2 where measurements for both stars could be obtained. Red is used for atomic lines, green for molecular lines, and blue for DIBs. All points show the EW uncorrected for saturation (which should be a good approximation in all cases except for the intense atomic lines) except for K I $\lambda 7664.911+7698.974$. In that case, besides the uncorrected values we show the corrected one (as an unfilled symbol) calculated using the kinematic decomposition described in the text. The x value in each case is λ_0 , the length of the horizontal error bars is proportional (with a threshold for the narrowest ones) to the FWHM of the line, and the vertical error bars show the uncertainty in the measurement. The two dotted lines are the ratios for $E(4405 - 5495)$ and A_V .

ing the same values as for LS III +46 11, pointing towards similar properties for both sightlines (as also indicated by the similar extinction values).

- The EW ratios between LS III +46 12 and LS III +46 11 are plotted in Fig. 3, along with the corresponding ratios for the dust measurements, 0.76 and 0.78 for $E(4405 - 5495)$ and A_V , respectively (as previously noted, the two stars show similar values of R_{5495} , hence the similar values for the two ratios).
- The behavior for the EW ratios of the molecular lines in Fig. 3 is very different. CH+ $\lambda 4232.548$ has a value of 1.09, clearly above the values for dust and similar to the values for the DIBs (see below). On the other hand, the value for CH $\lambda 4300.313$, 0.55, is clearly below the dust values and the one for CN $\lambda 3873.999+74.607+75.760$, 0.10, is much lower.
- Based on the previous points, if we model the ISM towards Berkeley 90 as a common cloud (cloud σ) that affects LS III +46 11 and LS III +46 12 similarly and a second cloud (cloud ζ) that affects LS III +46 11 almost exclusively³, cloud ζ would have little or nothing CH+ but a significant amount of dust, an even larger (in relative terms)

amount of CH, and more CN by an order of magnitude⁴. Given that CH traces gas denser than CH+ (Smoker et al. 2014), cloud ζ should be denser than cloud σ .

- A similar comparison for atomic EW ratios is less straightforward because of saturation effects associated with the larger optical depths. The ratio for K I $\lambda 7698.974$, which is optically thinner than the Na I and Ca II lines, is 0.85, which is relatively close to the dust ratios. See below for a more detailed analysis taking into account the kinematics and saturation.
- There is a significant scatter in the vertical axis of Fig. 3 for the DIBs, as expected, since the correlations between DIBs and extinction are good but not perfect. However, what is more surprising is that for all cases the EW ratio is higher than for $E(4405 - 5495)$ or A_V . If we follow the two-cloud hypothesis previously described, that would mean that cloud ζ is depleted in DIBs with respect to cloud σ .
- Within the scatter shown in Fig. 3, ζ -type DIBs (e.g. DIB $\lambda 5797.06$ and DIB $\lambda 5849.81$, Krelowski et al. 1997) have values of ≈ 1.0 or below while σ -type DIBs (e.g. DIB $\lambda 5780.48$) have values of ≈ 1.1 or above. This is another indication that cloud ζ should be interpreted as being denser than cloud σ .
- McCall et al. (2010) discovered that the 6195.98 Å and 6613.62 Å DIBs are nearly perfectly correlated. LS III +46 11 and LS III +46 12 are more extinguished than the majority of stars in their sample but the strong correlation still appears to hold. The EW ratios for those two lines in Fig. 3 are 1.037 ± 0.020 and 1.012 ± 0.014 i.e. one sigma away from each other, and the ratios of one line to another are 4.1–4.2, very close to the 3.96 value determined by McCall et al. (2010). Note, however, that a nearly perfect correlation does not necessarily imply a same carrier (Oka et al. 2013).

Second, we qualitatively discuss the existence of different kinematic components seen in some of the ISM lines:

- There are two clear kinematic components in K I $\lambda 7664.911+7698.974$ for LS III +46 11 and LS III +46 12 in the CAFÉ-BEANS data. A stronger one at -19.34 ± 0.01 km/s for LS III +46 11 (-20.96 ± 0.02 km/s for LS III +46 12) and a weaker one at -8.00 ± 0.02 for LS III +46 11 (-6.43 ± 0.04 km/s for LS III +46 12). The larger velocity difference between components for LS III +46 12 manifests itself in the height of the central peak in the top two spectra of Fig. 5.
- The strong component has a velocity consistent with that of LS III +46 11 measured in Paper I and is relatively close to that of LS III +46 12. The weak component, on the other hand, has a velocity intermediate between those of the stars and that of the Sun. We adopt as our model that the strong component originates in a cloud associated with Berkeley 90 and that the weak one originates in a cloud in the path between the cluster and the Sun. In terms of the σ and ζ clouds defined before, the strong component would contribute to both clouds σ and ζ while the weak one would contribute only to σ . Or, putting it in another way, the ISM common to both sightlines could be described by a cloud σ_1 far away from the cluster (the weak kinematic component) and a cloud

⁴ We should think of these two clouds not as two exclusive physical entities. Indeed, as described below, cloud ζ appears to be associated with the cluster while cloud σ appears to have two components, one associated with the cluster and one located in the path between Berkeley 90 and the Sun.

³ The reason for the cloud nomenclature is explained later on.

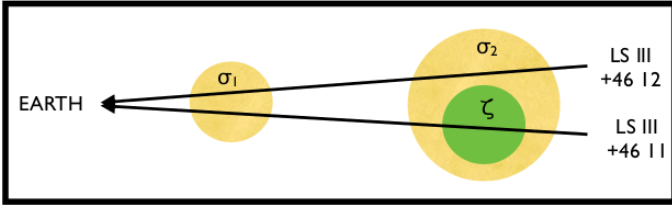


Fig. 4. Model for the ISM clouds present in the LS III +46 11 and LS III +46 12 sightlines. Cloud σ_1 is of low density, not associated with the cluster, and affects both sightlines similarly. Cloud σ_2 is also of low density, is the skin of the cloud associated with the cluster, and is longer along the LS III +46 12 sightline. Cloud ζ is of high density, is the core of the cloud associated with the cluster, and affects LS III +46 11 exclusively (or, at least, to a much larger degree than LS III +46 12).

σ_2 associated with the cluster (part of the strong kinematic component). In addition, the LS III +46 11 sightline would have another (denser) cloud ζ associated with the cluster (i.e. with the same velocity as σ_2). In the most simplified version of the model, σ_1 and σ_2 would produce the same effect in the spectra of LS III +46 11 and LS III +46 12 but we know that this cannot be completely true, as some absorption lines have values above 1.0 in Table 3. This can be explained by σ_2 being longer along the LS III +46 12 sightline than along the LS III +46 11 one. Figure 4 represents this description.

- Na I $\lambda 5889.951+5895.924$ is saturated for both stars but the profiles are consistent with the existence of the two same kinematic components as for K I.
- Ca II $\lambda 3933.663+3968.468$ for LS III +46 11 is also close to saturation and the only spectra with good S/N are from NoMaDS, so the resolution is lower than for K I or Na I. Also, for LS III +46 12 the high-resolution spectra have low S/N in the 4000 Å region so we only have GOSSS data there and only for Ca II $\lambda 3933.663$, since it is not possible to separate Ca II $\lambda 3968.468$ from H ϵ at the GOSSS lower resolution. In any case, the available data are consistent with the existence of the two kinematic components detected in the K I lines.
- The existence of two kinematic components can also be seen in the asymmetric profile of CH+ $\lambda 4232.548$ and CH $\lambda 4300.313$ in LS III +46 11.
- We do not detect two kinematic components in CN $\lambda 3873.999+3874.607+3875.760$, CH $\lambda 3886.410$, CH+ $\lambda 3957.692$, or Ca I $\lambda 4226.7275$, quite possibly due to an insufficient S/N. However, their measured velocities in Table 2 indicate that they originate mostly in the strong K I kinematic component.
- In addition, there is a third high-velocity redshifted component detectable only for LS III +46 12 in Na I $\lambda 5889.951+5895.924$ and Ca II $\lambda 3968.468$. The component is weak in Na I but the high-resolution data allows us to measure a precise velocity of $+196.04 \pm 0.22$ km/s. The third component is much stronger in Ca II $\lambda 3968.468$, making it easily detectable even at the GOSSS resolution (but note that we do not see it in either LS III +46 11 or LS III +46 11 B). The strength of Ca II $\lambda 3968.468$ is a common feature in high-velocity absorption components due to the lifting of the calcium depletion (Routly & Spitzer 1951; Walborn et al. 2002) and is the reason why Ca II $\lambda 3968.468$ appears near the top of Fig. 3. The high-velocity component is not seen in the K I lines.

- We were unable to detect the two kinematic components in any of the DIBs. This is expected, given the spectral resolution of our data, the small velocity difference between the two clouds, and the asymmetric and possibly variable profile of the narrow strong DIBs (Galazutdinov et al. 2008; Oka et al. 2013).

Third, we analyze the widths and column densities by kinematic component (Fig. 5 and Table 6) and calculate the column densities for the atomic and molecular species (whenever possible):

- The values for the b velocity parameter for the strong kinematic component are rather constant across different lines for the two stars, with values between 3.2 and 3.6 km/s. The values for the weak component are similar but with a larger scatter, possibly due to the fitting noise induced by the presence of a stronger component. Given the large values of b , its most likely origin is “turbulence” (i.e. multiple components not detectable with the available resolution) rather than thermal motions.
- The weak component has very similar K I column densities for both stars. That fact agrees with our model that it originates in a cloud (σ_1) not associated with Berkeley 90.
- When we compute the total column-density ratio between LS III +46 12 and LS III +46 11 for K I, we find a value of 0.76 ± 0.07 , which is in agreement (within the errors) with the ratios for $E(4405 - 5495)$ and A_V (Fig. 3). It appears that the lower value obtained directly from the EW ratios is indeed caused by a saturation effect. In other words, the comparison between the two sightlines provides a good correlation between the amounts of dust and K I which is better than for any of the measured DIBs.
- Even though the uncertainty for the CH+ $\lambda 4232.548$ weak component is large, the ratio between the column densities of CH and CH+ is larger for the strong component than for the weak one, pointing in favor of (some of) the strong component originating in a denser medium than the weak one. More specifically, if we include the values in the last panel of Fig. 5 of Smoker et al. (2014), [a] the weak component would have large values of $N(\text{CH})$ and $N(\text{CH}+)$ but they would follow the average relation while [b] the strong component would be located at the top right corner (very large values of both column densities) with a larger than average $N(\text{CH})/N(\text{CH}+)$.
- We can only compare the total column densities of CN with those of CH and CH+. For LS III +46 12 the three column densities fall reasonably well within the trends shown in Fig. 5 of Smoker et al. (2014). For LS III +46 11, CH and CN fall within their trend but both of them are above their expected values with respect to CH+. This indicates that all clouds (and especially cloud ζ) are “CN-like CH” i.e. dense. If we use the Weselak et al. (2008) sample as a comparison, both stars would be among the ones with the highest CH column densities. LS III +46 11 would also be in that category for CN but LS III +46 12 would be among the targets with low column densities.

5. Discussion

Our analysis of the ISM in front of Berkeley 90 points towards the existence of [a] one component (σ_1) located at a different velocity (and likely distance) than that of the cluster that affects both stars similarly and [b] a second component at a velocity similar to that of the cluster that affects LS III +46 11 and

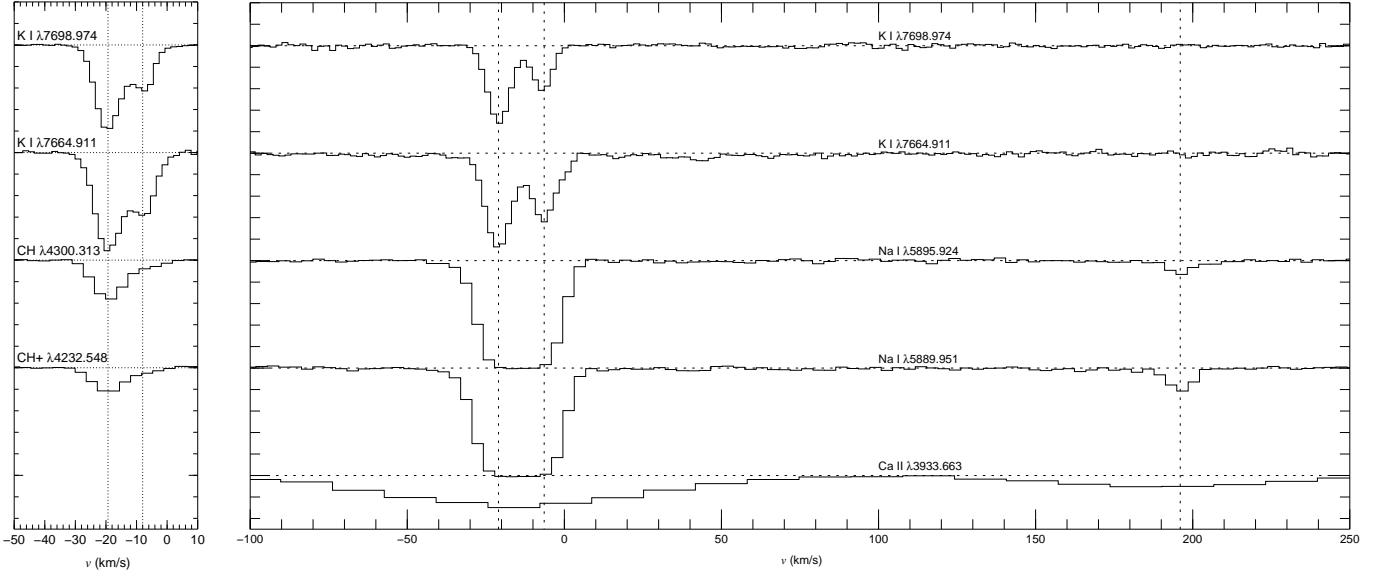


Fig. 5. Line profiles with multiple components for LS III +46 11 (left) and LS III +46 12 (right). The K I and Na I spectra are from CAFÉ-BEANS, the CH and CH+ spectra from NoMaDS, and the Ca II from GOSSS (hence, the significantly lower resolution, though the final spectrum has been drizzled from the original data to improve it). The horizontal dotted lines indicate the normalized continuum level (the spectra are separated by one unit) and the vertical dotted lines indicate the velocities measured for K I and for the third high-velocity component.

Line(s)	LS III +46 11			LS III +46 12		
	St. comp.	Wk. comp.	Total	St. comp.	Wk. comp.	Total
CN $\lambda 3873.999+74.607+75.760$			12.98 ± 0.05			11.99 ± 0.06
CH+ $\lambda 3957.692$			13.79 ± 0.03			13.83 ± 0.07
Ca I $\lambda 4226.7275$			10.70 ± 0.04			10.76 ± 0.08
CH+ $\lambda 4232.548$	13.73 ± 0.02	13.25 ± 0.20	13.85 ± 0.02			13.89 ± 0.05
CH $\lambda 4300.313$	13.99 ± 0.02	13.19 ± 0.08	14.05 ± 0.02			13.74 ± 0.05
K I $\lambda 7664.911+7698.974$	12.40 ± 0.01	11.84 ± 0.04	12.50 ± 0.02	12.23 ± 0.01	11.85 ± 0.03	12.38 ± 0.02

Table 6. Logarithm of the measured ISM column densities (in cm^{-2}). When high-resolution spectra are available, the column densities for the two components plus the total value are given using the kinematic decomposition described in the text. In the other cases, we used the EWs from Tables 2 and 4. Since for CH+ $\lambda 4232.548$ and CH $\lambda 4300.313$ no high-resolution spectra were available for LS III +46 12 and those lines have a slight saturation effect, we applied a correction to the EW based on the comparison between the intermediate-resolution and high-resolution results for those lines in LS III +46 11. The molecular oscillator strengths used are those of Smoker et al. (2014) while the ones for the atomic lines were obtained from the NIST web page <http://physics.nist.gov/Pubs/AtSpec/table105.html>.

LS III +46 11 B (as $\sigma_2 + \zeta$) more than LS III +46 12 (as just σ_2). All clouds appear to be relatively dense but cloud ζ is denser than the average of clouds σ_1 and σ_2 . This is shown by several indicators when comparing the equivalent widths different absorption lines for LS III +46 11 (the star with the larger extinction) and LS III +46 12:

- The EWs for CH $\lambda 4300.313$ and CN $\lambda 3873.999+74.607+75.760$ increase more than the extinction while CH+ $\lambda 4232.548$ increases less.
- K I $\lambda 7664.911+7698.974$ increases in a similar way as extinction (once saturation is taken into account).
- All DIBs increase less than extinction (the skin effect, Herbig 1995).
- ζ -type DIBs (e.g. DIB $\lambda 5797.06$) increase more than σ -type DIBs (e.g. DIB $\lambda 5780.48$). ζ -type DIBs originate in different parts of the ISM, including UV-shielded, dense regions, while σ -type DIBs originate preferentially in UV-exposed, thin regions of the ISM and are strongly depleted in dense

regions (Krelowski et al. 1997; Cami et al. 1997). This DIB property is the reason for the chosen cloud nomenclature.

- For the limited amount of lines with information for LS III +46 11 B, the ISM in front of that star appears to be more similar to LS III +46 11 (to which is spatially closer) than to LS III +46 12.

Therefore, we predict that the remains of the molecular cloud from which Berkeley 90 formed are predominantly located towards the NW. This description is also consistent with the WISE images, which show that the cluster has not evacuated a cavity yet and that there is an irregular dust distribution throughout the cluster. Hence, the surrounding ISM also points towards a young age for the cluster.

As previously mentioned, there are few sightlines in the literature with a large number of DIBs and even fewer for targets with extinction as high as these and with the ISM lines measured for the whole optical spectrum. The additional interest in this work is that we have a pair of sightlines with a common absorption component and another which affects only one star,

#	λ_0 (Å)	EW ratio	FWHM (Å)
1	4501.67	0.822±0.047	3.2
2	6590.42	0.832±0.068	12.8
3	4726.70	0.866±0.015	3.9
4	5418.87	0.882±0.168	0.7
5	4761.12	0.882±0.059	22.0
6	6113.18	0.885±0.058	0.7
7	6379.32	0.890±0.057	0.6
8	6699.32	0.892±0.065	0.6
9	5797.06	0.932±0.021	0.8
10	6439.48	0.956±0.069	0.8
11	4779.69	0.957±0.078	3.0
12	5449.83	0.959±0.054	14.1
13	5849.81	0.960±0.027	0.8
14	6425.66	0.988±0.079	0.8
15	6203.05	1.025±0.034	1.2
16	6613.62	1.037±0.025	0.9
17	6195.98	1.039±0.020	0.4
18	6993.13	1.052±0.035	0.8
19	4879.83	1.054±0.113	11.5
20	5494.29	1.067±0.075	0.7
21	6672.27	1.094±0.093	0.7
22	4427.94	1.094±0.020	17.3
23	6597.31	1.097±0.173	0.5
24	6660.71	1.099±0.039	0.6
25	6269.85	1.120±0.026	1.2
26	7224.03	1.121±0.016	1.0
27	6010.75	1.132±0.037	3.3
28	4762.36	1.134±0.071	2.7
29	5780.48	1.134±0.008	2.1
30	6283.84	1.140±0.029	4.8
31	6456.01	1.142±0.099	0.9
32	5705.08	1.152±0.037	3.7
33	6177.30	1.155±0.049	23.1
34	6445.28	1.166±0.078	0.6
35	5487.23	1.193±0.082	4.8
36	8621.20	1.194±0.083	4.7
37	4887.43	1.288±0.181	39.8
38	5404.56	1.362±0.198	0.8

Table 7. Sorted DIB equivalent-width ratios between LS III +46 12 and LS III +46 11 for the cases in Table 2 where measurements for both stars could be obtained. DIBs near the top should be less depleted in ζ clouds (dense), the prototype of such cases being DIB $\lambda 5797.06$. DIBs near the bottom should be more depleted in dense clouds and more prominent in σ clouds (diffuse), the prototype of such cases being DIB $\lambda 5780.48$. Note that the uncertainties in some equivalent-width ratios are relatively high, so their placement in the sequence is uncertain. The FWHM is shown to differentiate narrow and broad DIBs.

with the second one being [a] dense and [b] DIB-depleted. This provides the opportunity to rank the measured DIBs in a $\sigma - \zeta$ scale (Krelowski et al. 1997; Cox et al. 2005) by sorting them by the ratios of the equivalent widths plotted in Fig. 3. We have done that in Table 7:

- The two DIBs that combine the properties of being apparently less affected by depletion in the dense ISM with being relatively narrow and strong are DIB $\lambda 4501.67$ and DIB $\lambda 4726.70$. Those would be good candidates to correlate better with extinction even when the ISM becomes very dense, a hypothesis that we plan to test in the future with other sightlines with high extinction (note that previous studies that included DIB $\lambda 4726.70$ such as Puspitarini et al.

(2013) tend to probe just the diffuse ISM). At longer wavelengths the best candidate is still the ζ -prototype, DIB $\lambda 5797.06$.

- At the other end of the scale, DIB $\lambda 8621.20$, present in the Gaia band, is a good example of a strong DIB that is expected to correlate poorly with extinction because it appears to be highly depleted in the dense ISM. A similar case is DIB $\lambda 5487.23$, which already had the worst correlation with extinction of the eight DIBs in Friedman et al. (2011).
- An interesting result is that for the pair DIB $\lambda 5404.56$ + DIB $\lambda 5418.87$. They have relatively large errors in Table 7 but they are close in wavelength and in opposite sides of the $\sigma - \zeta$ scale, making them good candidates to sample the diverse behavior of DIBs with density. Unfortunately, they are not very strong and they are placed at the wings of He II $\lambda 5411.53$, which is strong for O stars.
- The order in Table 7 is in good agreement with the correlation with extinction for eight DIBs of Friedman et al. (2011).
- There is no apparent correlation between the $\sigma - \zeta$ scale and the width of the DIBs: there are both narrow and broad DIBs at both ends of the scale.

Another interesting aspect of the intervening ISM is the location of Berkeley 90 well above the Galactic Plane. With a $\log d$ of 3.40–3.45 and considering the Sun’s own vertical distance (Maíz Apellániz 2001; Maíz Apellániz et al. 2008), we find that Berkeley 90 is located almost 200 pc above the Galactic Plane, or about six times the scale height for OB stars. Even for a $\log d \sim 3.2$, Berkeley 90 would be ~ 150 pc above the Galactic Plane.

We have convincingly shown that the amount of extinction is variable across the face of Berkeley 90, with the value for LS III +46 11 being greater than that of LS III +46 12 by 30%, even though they are separated by a little over a pc in the plane of the sky. Also, the additional extinction experienced by LS III +46 11 appears to take place in a relatively dense cloud with properties different to the average of the ones that cause the rest of the extinction along the line of sight (regarding e.g. DIBs). Yet, despite those differences, the measured R_{5495} for both stars is essentially the same, indicating that there are no large differences in the average dust grain size for both sightlines. In other words, environments with different dust grain sizes are not required to produce differences in the observed DIBs.

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